Spectral evolution of flaring blazars from numerical simulations

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Outline

- Motivation
- Simulation Setup
- Results
- Summary and Outlook

Motivation

2006 radio flare in CTA 102 (z=1.037)



- single dish data 4 GHz -345 GHz (UMRAO, Metsähovi, SMA)
- VLBA observations at 15GHz and 43GHz (MOJAVE & Boston)
- VLBA multi-freq. observations 2GHz -86 GHz

Light curve analysis



Ref: Marscher & Gear (1985), Fromm et al. (2011)

VLBI observations of CTA102

15GHz Monitoring

(macro-physics)

Multi-frequency (micro-physics)





Ref: Lister et al. 2009, Fromm et al. 2013a, 2013b

Results of the observations

several stationary features —>recollimation shocks first at 18pc (43GHz,86GHz)

interactions between travelling and recollimation shock interactions



simulations to investigate shock-shock interactions (focusing on the first recollimation shock)

relativistic hydrodynamics+emission simulations(Perucho et al. 2010)(Mimica et al. 2009 & Fromm 2013)

RHD simulations

over-pressured jet



pressure matched jet





9.86e-03

1.01e-03

1.03e-04









Radiative transfer (ray tracing)

include micro-physics: absorption, emission, losses and delays

using 3D ray-tracing technique and large frequency range



synchrotron spectrum (1–1000GHz) at each pixel

Single dish light curves



Radio maps

over-pressured jet



pressure matched jet



Shock-Shock interaction (re-scaled)

eRHD simulations

CTA102@15GHz





Summary & Outlook

- recollimation shocks in the jet of CTA102
- simulations of recollimation shocks
- spectral evolution of shocks in jets
- fake observations —> observable signatures

★ RMHD simulations of jets
 ★ polarised radiative transfer
 ★ 3D R(M)HD simulations (test stability of RCS)

Emission Simulation

$$\begin{split} n\left(\gamma\right) &= n\left(\gamma_{\min}\right) \left(\frac{\gamma}{\gamma_{\min}}\right)^{-p} \quad \gamma_{\min} < \gamma < \gamma_{\max} \qquad \text{e- distribution} \\ B &= \left(\overline{\epsilon_b} \frac{8\pi p_{th}}{\widehat{\gamma} - 1}\right)^{1/2} \qquad \qquad 0 < \epsilon_e < 1 \qquad \text{magnetic field [G]} \\ \gamma_{\max} &= \left(\frac{9m_e^2c^4}{8\pi e^3\epsilon_aB}\right)^{1/2} \qquad \qquad 0 < \epsilon_e < 1 \qquad \text{magnetic field [G]} \\ \eta_{\max} &= \left(\frac{9m_e^2c^4}{8\pi e^3\epsilon_aB}\right)^{1/2} \qquad \qquad 1e3 < \epsilon_a < 1e6 \qquad \text{max e- Lorentz factor} \\ \gamma_{\min} &= \frac{\overline{\epsilon_e}p_{th}m_p(p-2)}{\rho_{th}(\widehat{y} - 1)m_ec^2(p-1)} \qquad \qquad \text{min e- Lorentz factor} \\ n\left(\gamma_{\min}\right) &= \frac{\overline{\epsilon_e}p_{th}(p-2)}{(\widehat{\gamma} - 1)\gamma_{\min}^2m_ec^2} \left[1 - \left(\frac{\gamma_{\max}}{\gamma_{\min}}\right)^{2-p}\right]^{-1} \qquad \text{coeff. e- distribution} \end{split}$$

Ref: Mimica et al. (2009, 2010), Dermer & Boettcher (2010)

Emission Simulation

evolution of e- Lorentz factor (see Mimica et al. 2009)



e-Lorentz factor:

$$\gamma(\sigma) = \gamma_0 \frac{k_a e^{k_a \Delta \sigma}}{k_a + \gamma_0 k_s \left(e^{k_a \Delta \sigma} - 1\right)}$$

 $1 k \Lambda \sigma$

coeff. e- distr.

$$n_0(\gamma(\sigma)) = n_0(\gamma_0) \left[e^{k_a \Delta_\sigma} \left(1 + \gamma_0 \frac{k_s}{k_a} \left(e^{k_a \Delta \sigma} - 1 \right) \right) \right]^2$$

43 GHz evolution



Single dish light curves

over-pressured jet

pressure matched jet



Single dish analysis



Ref: Fromm et al. 2016

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Large scale simulations

Log Rest-mass density







using VLBA properties and observation settings



Radio maps (back-up)

over-pressured jet



Radio maps (back-up)

pressure-matched jet

